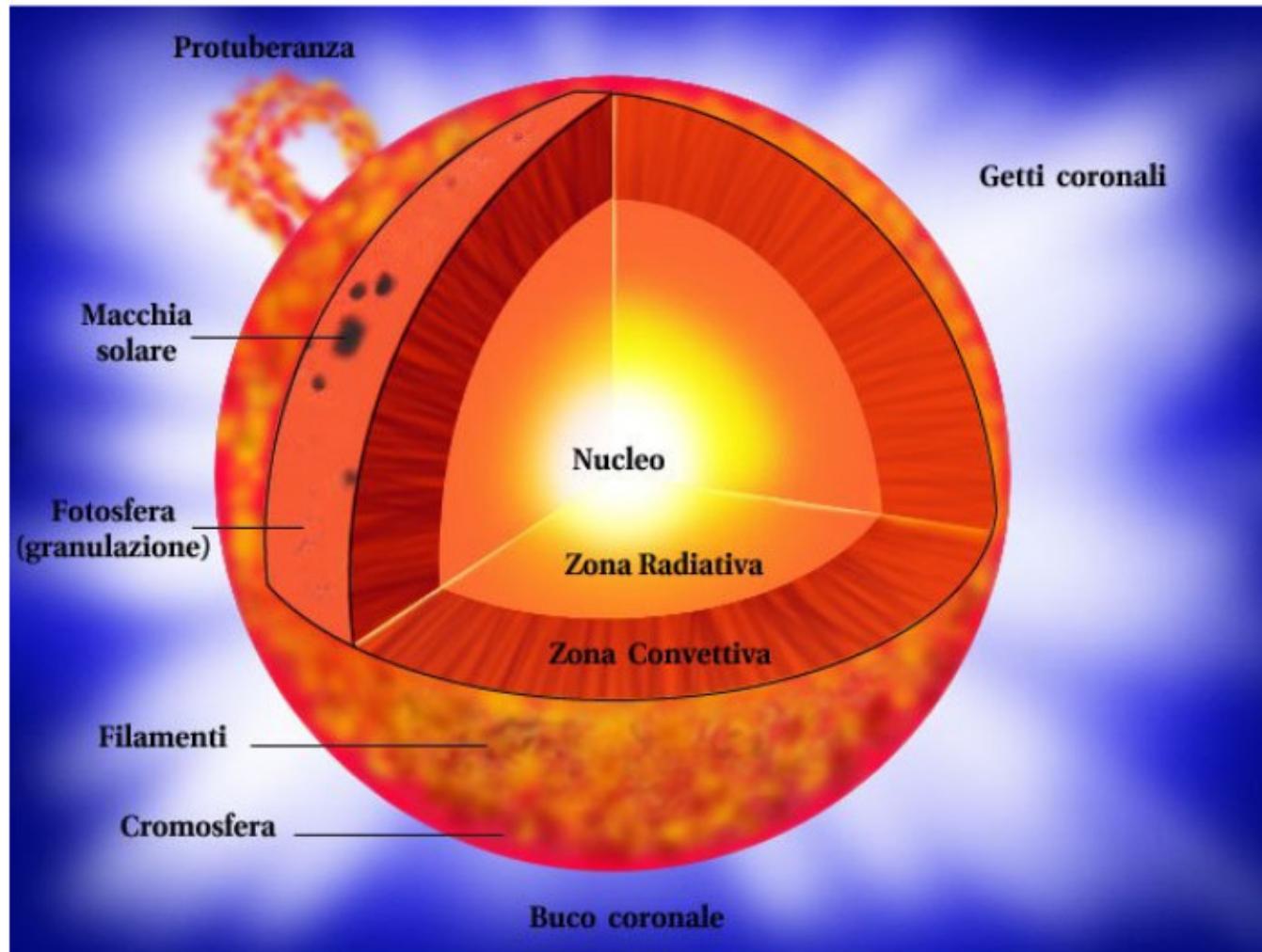
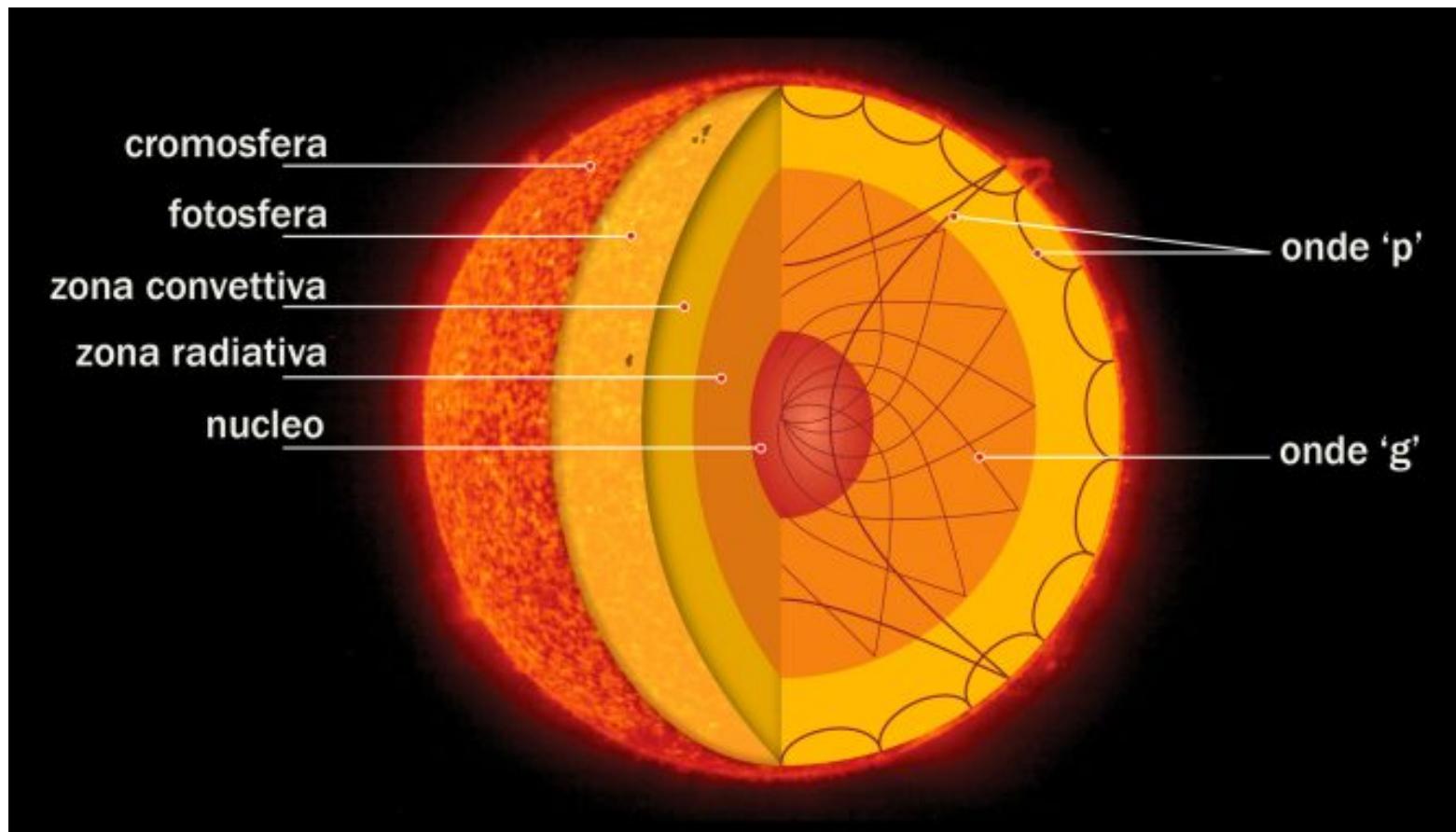


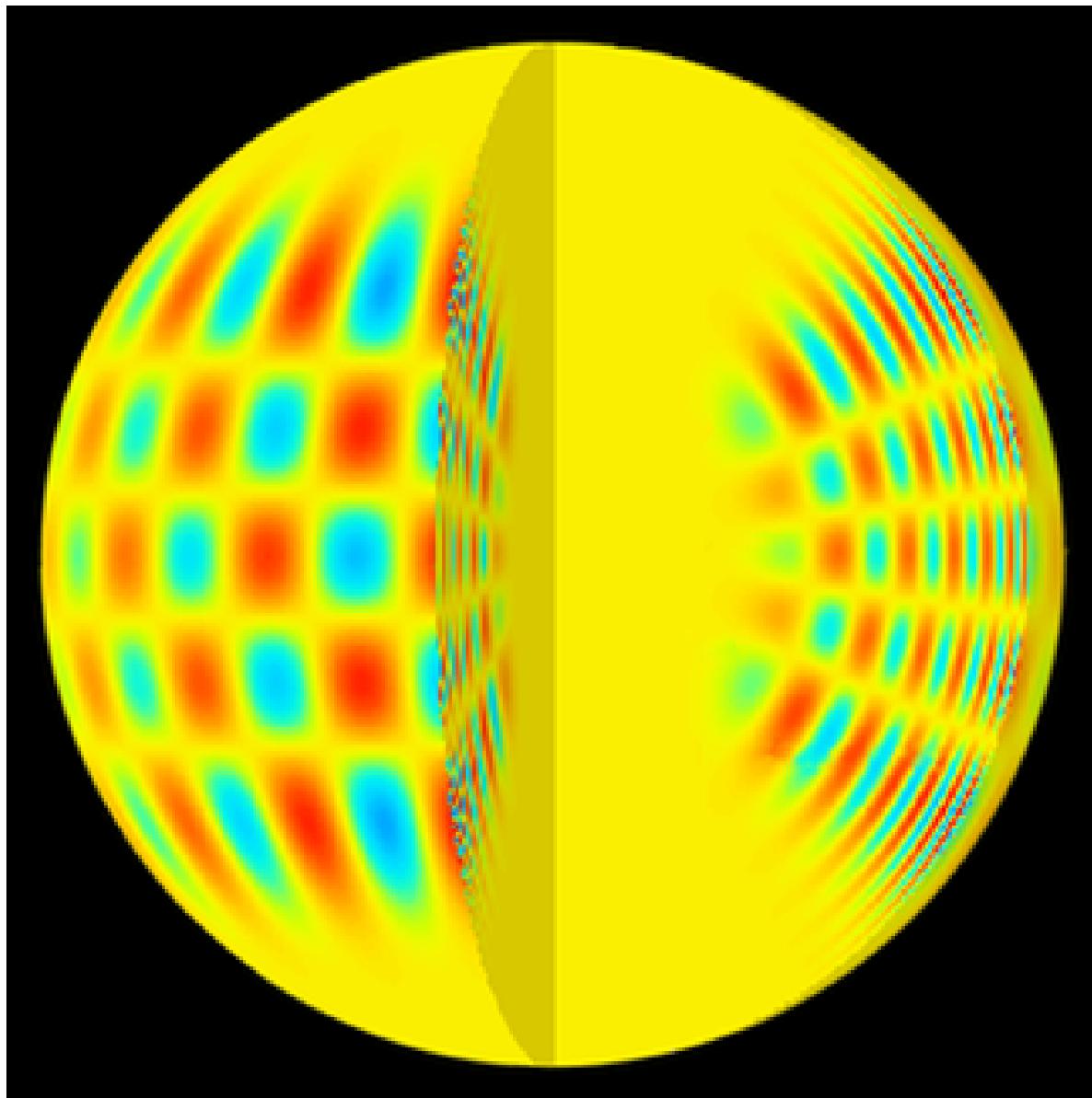
La struttura interna del Sole



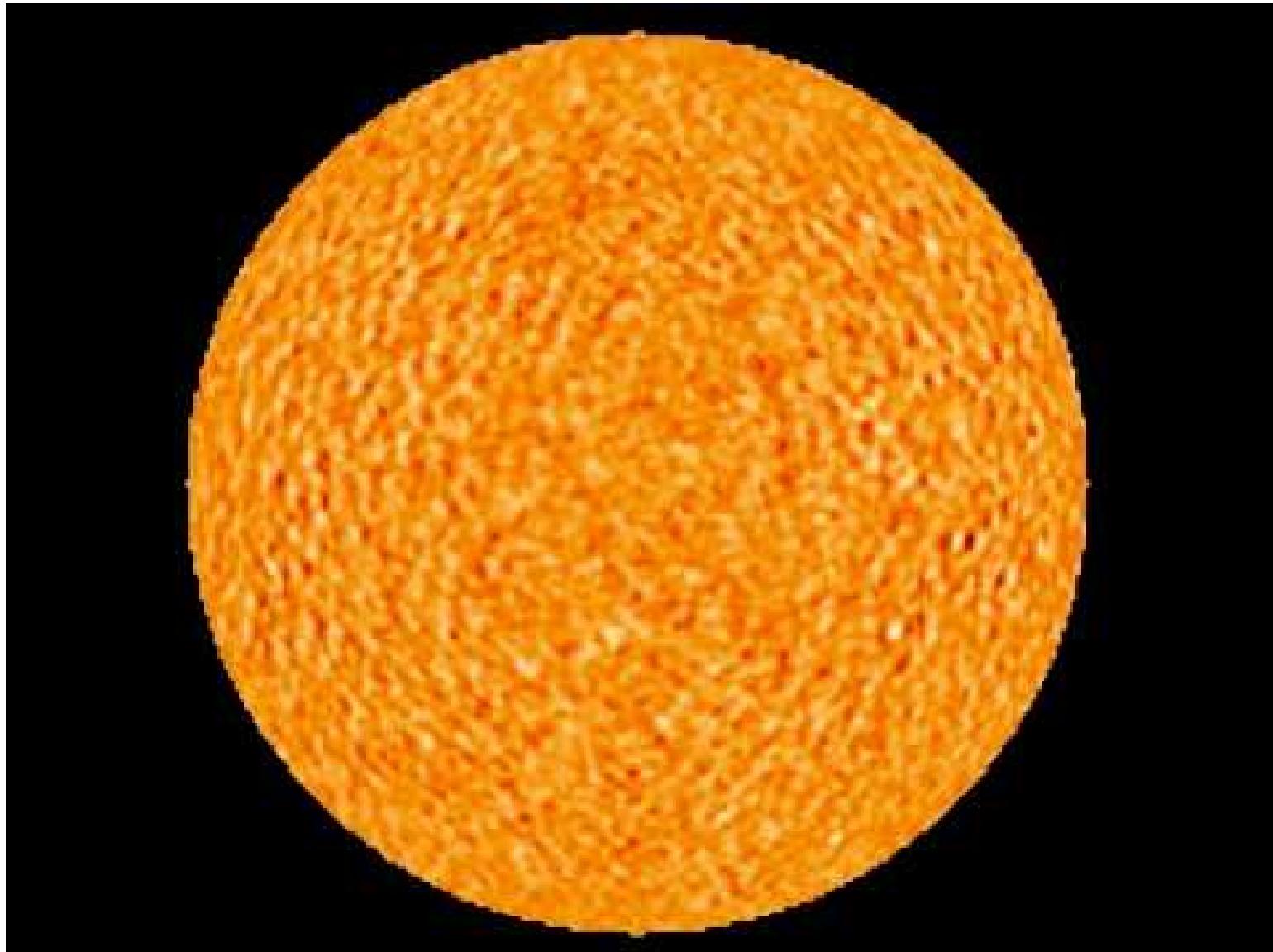
Eliosismologia



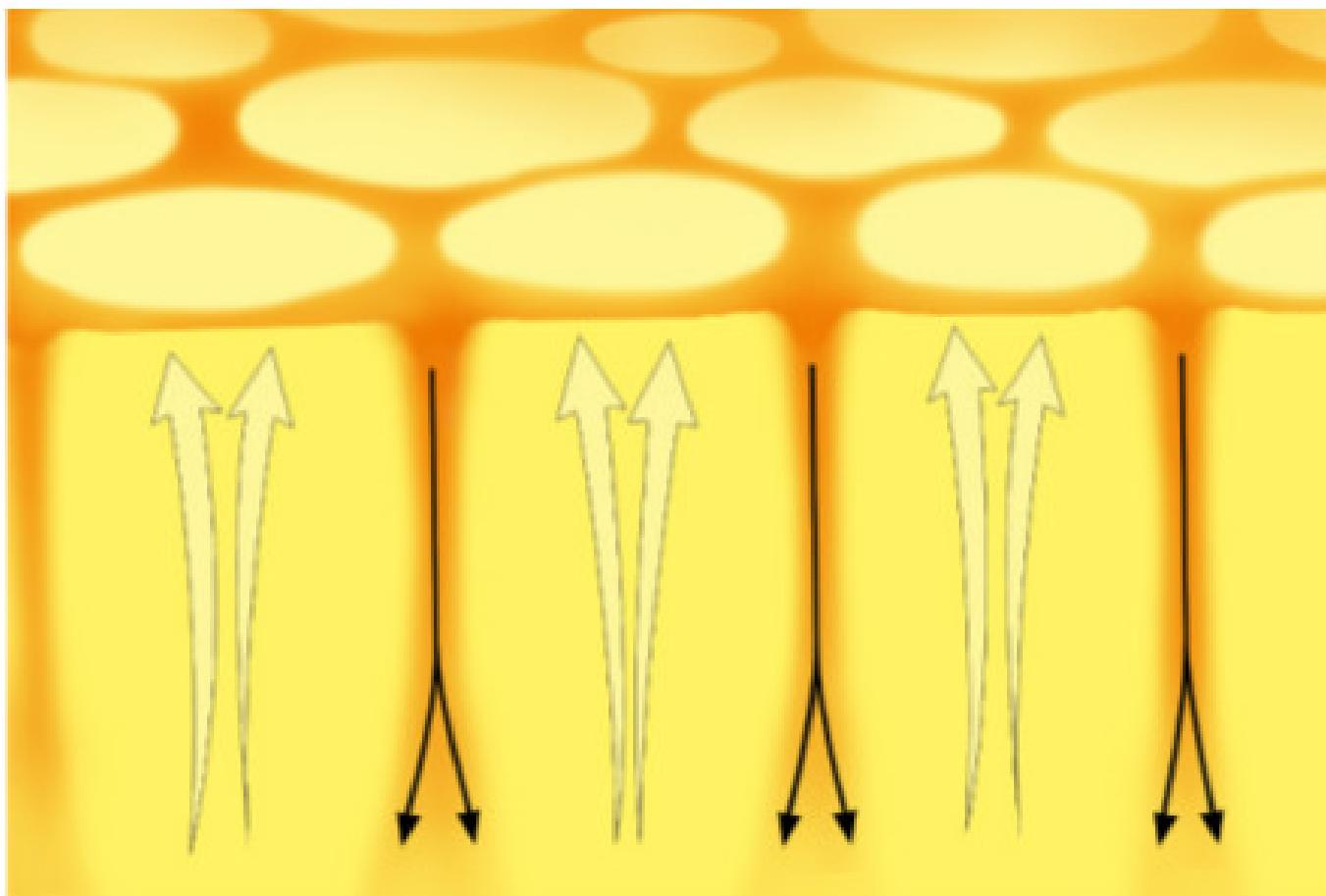
Onde stazionarie



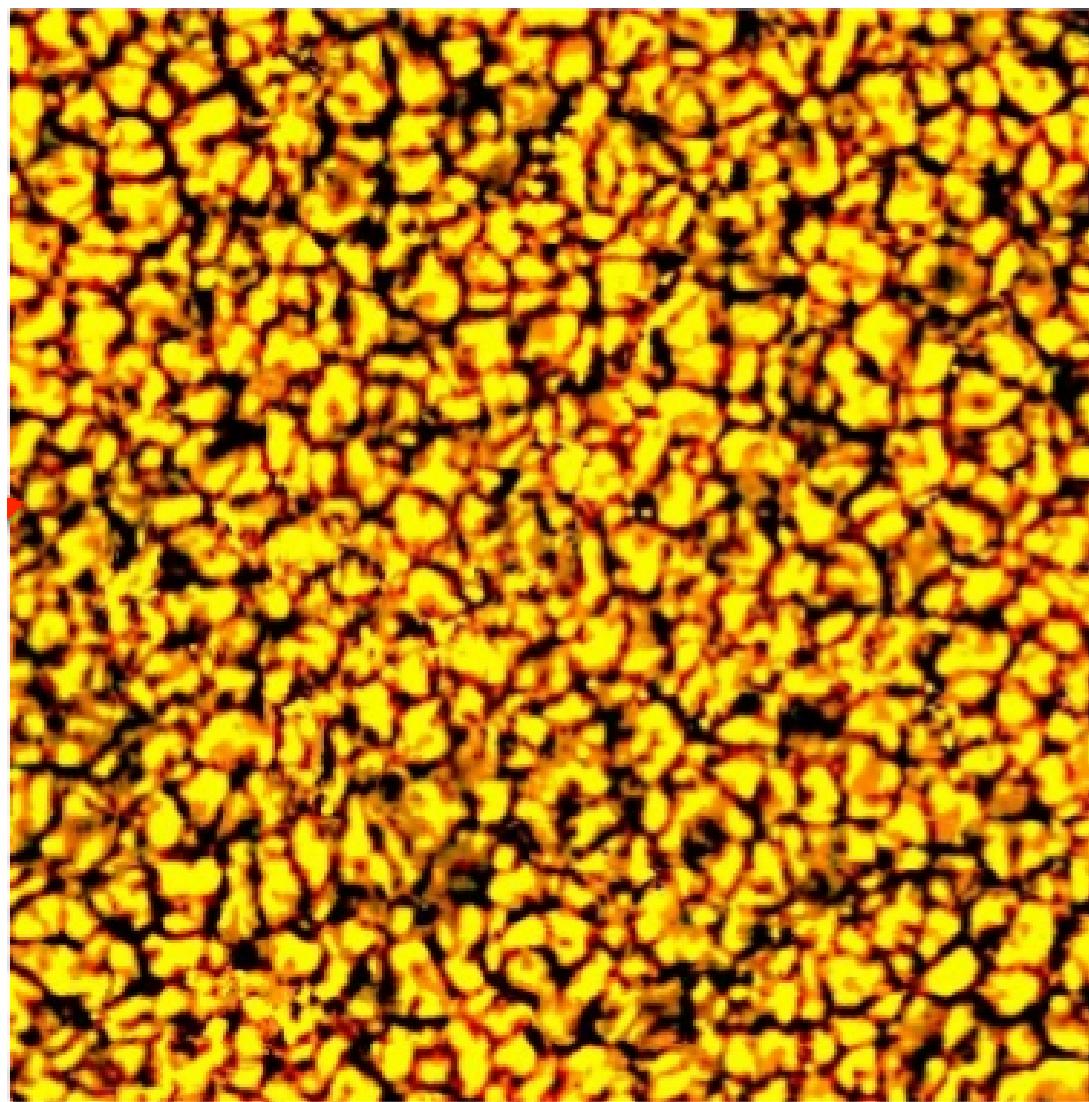
Modi osservati: effetto Doppler



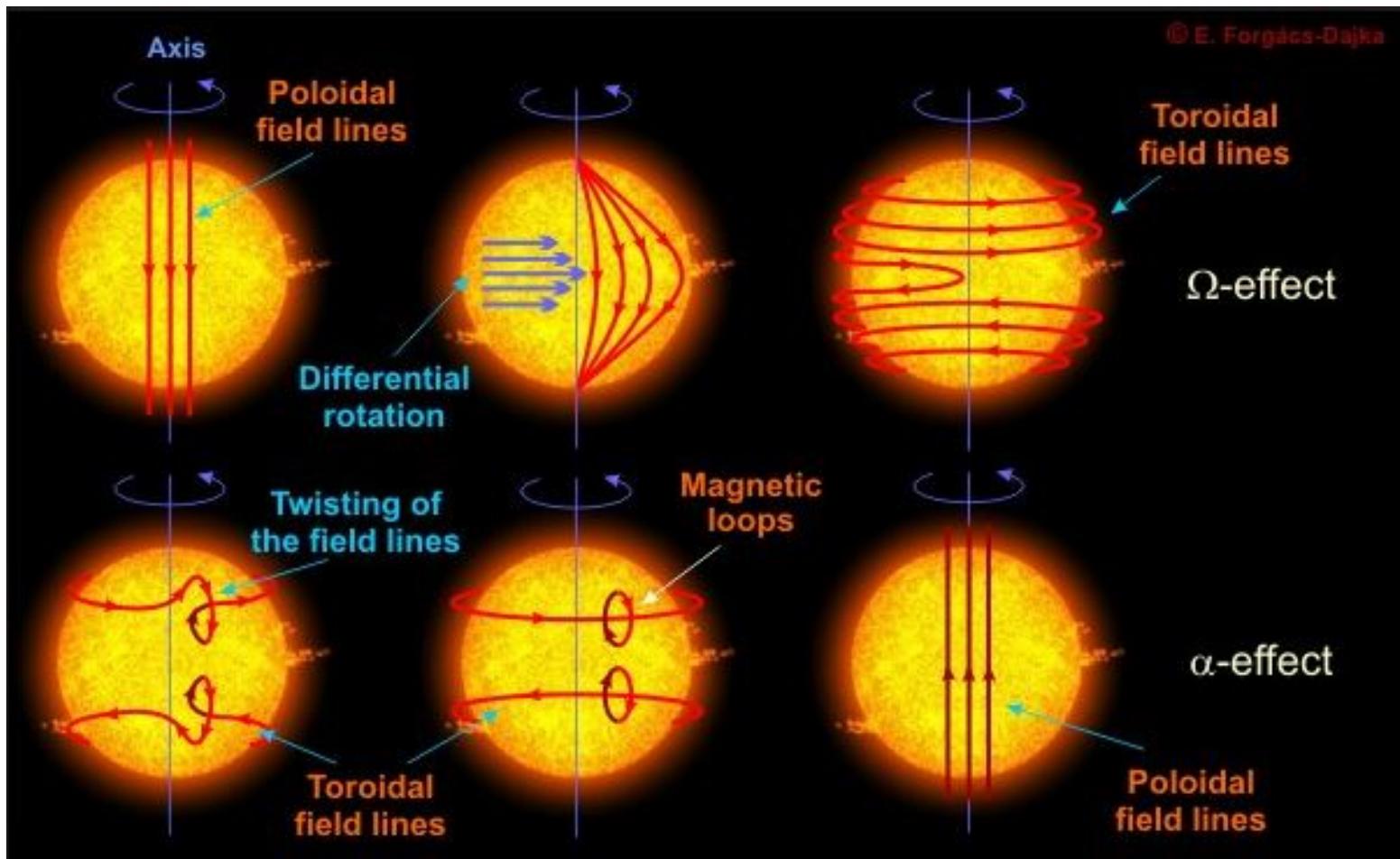
Celle convettive



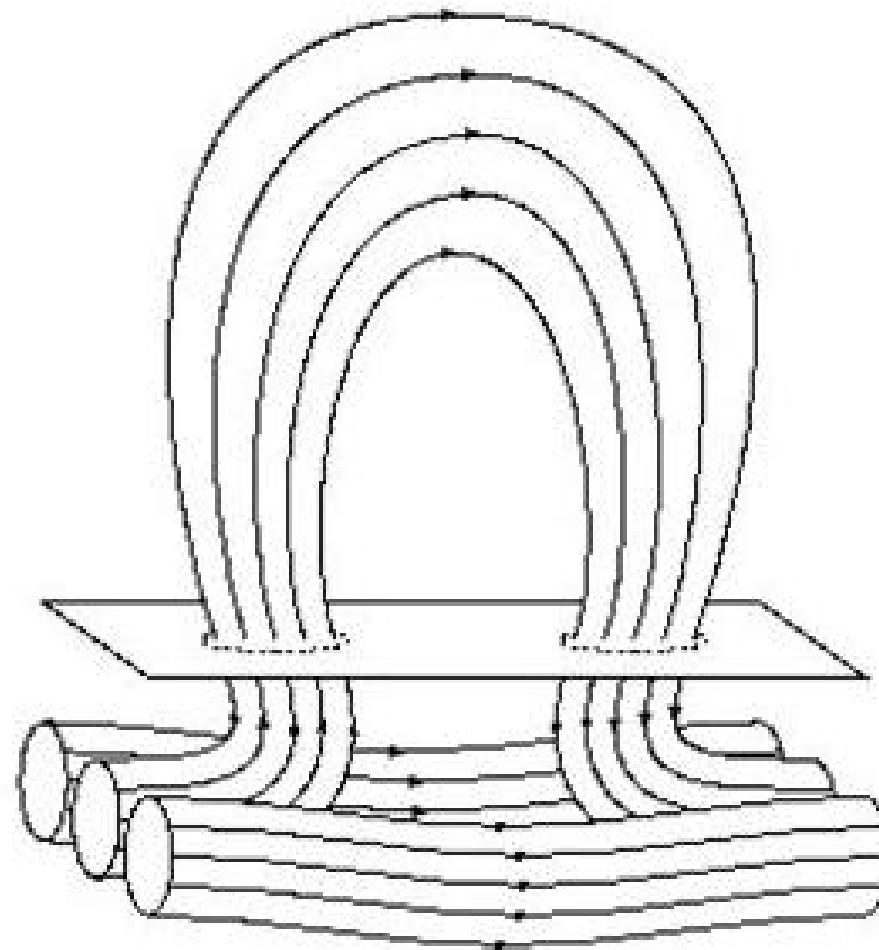
La granulazione



L'effetto dinamo nel Sole



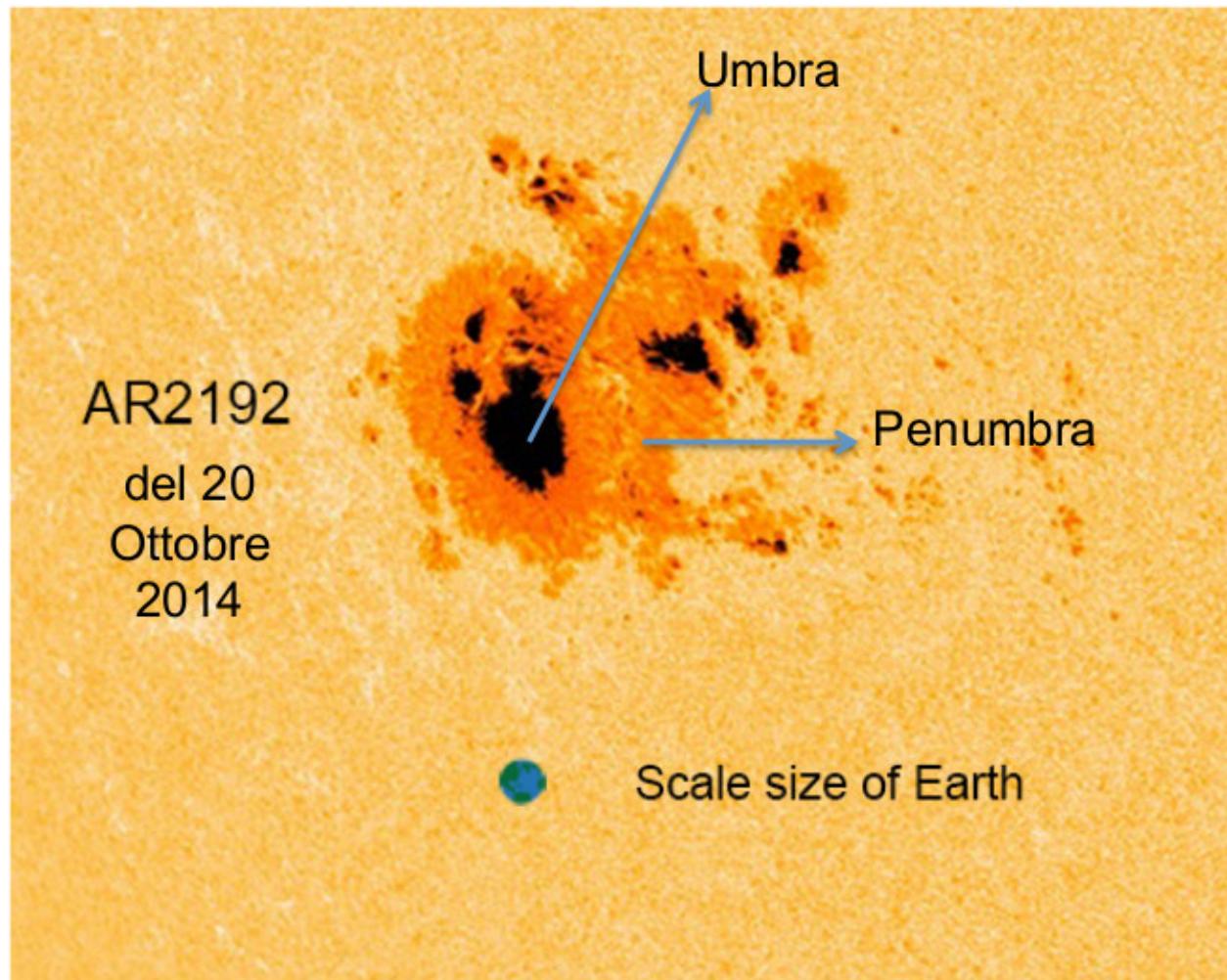
Galleggiamento magnetico



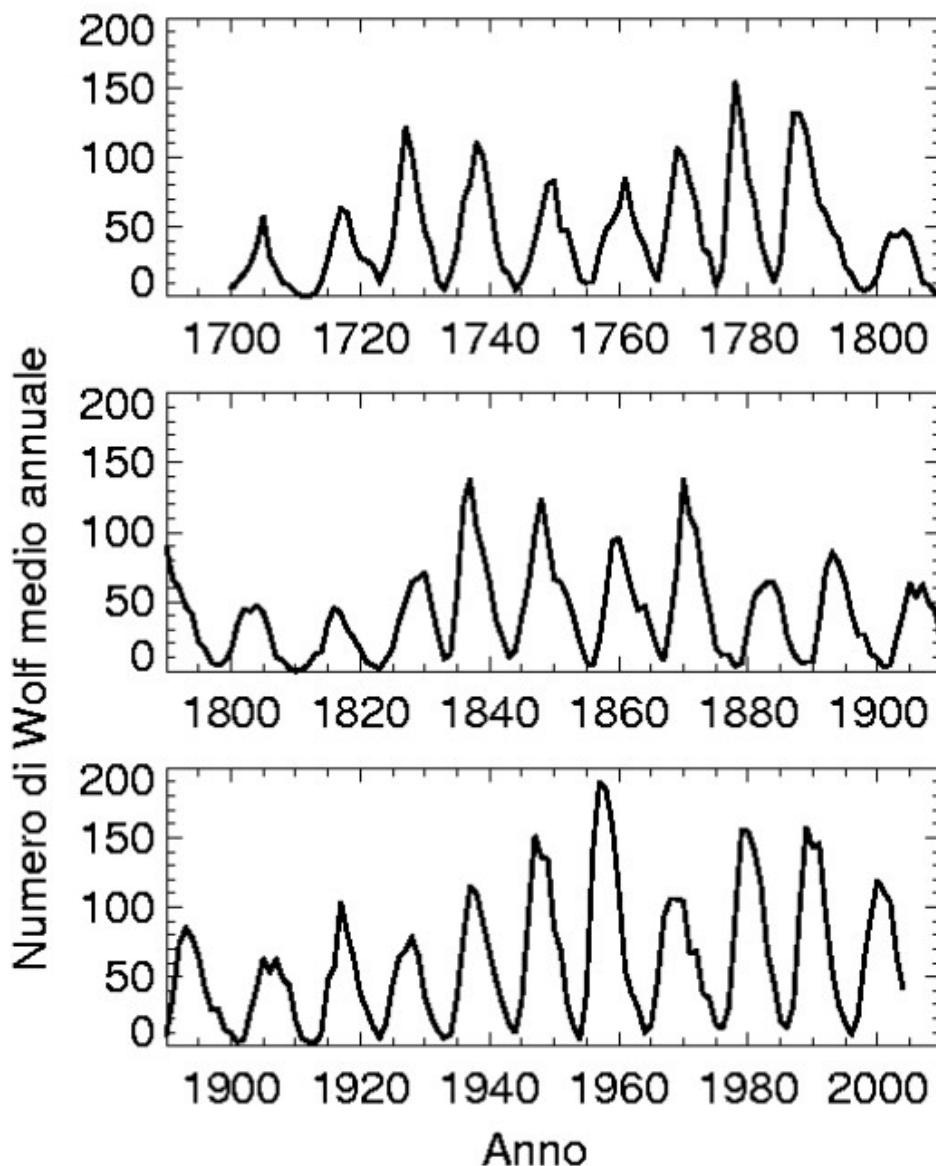
Macchie solari



La struttura delle macchie



Ciclo dell'attività solare



$$\text{Numero di Wolf} = k (m + 10g)$$

m = # di macchie

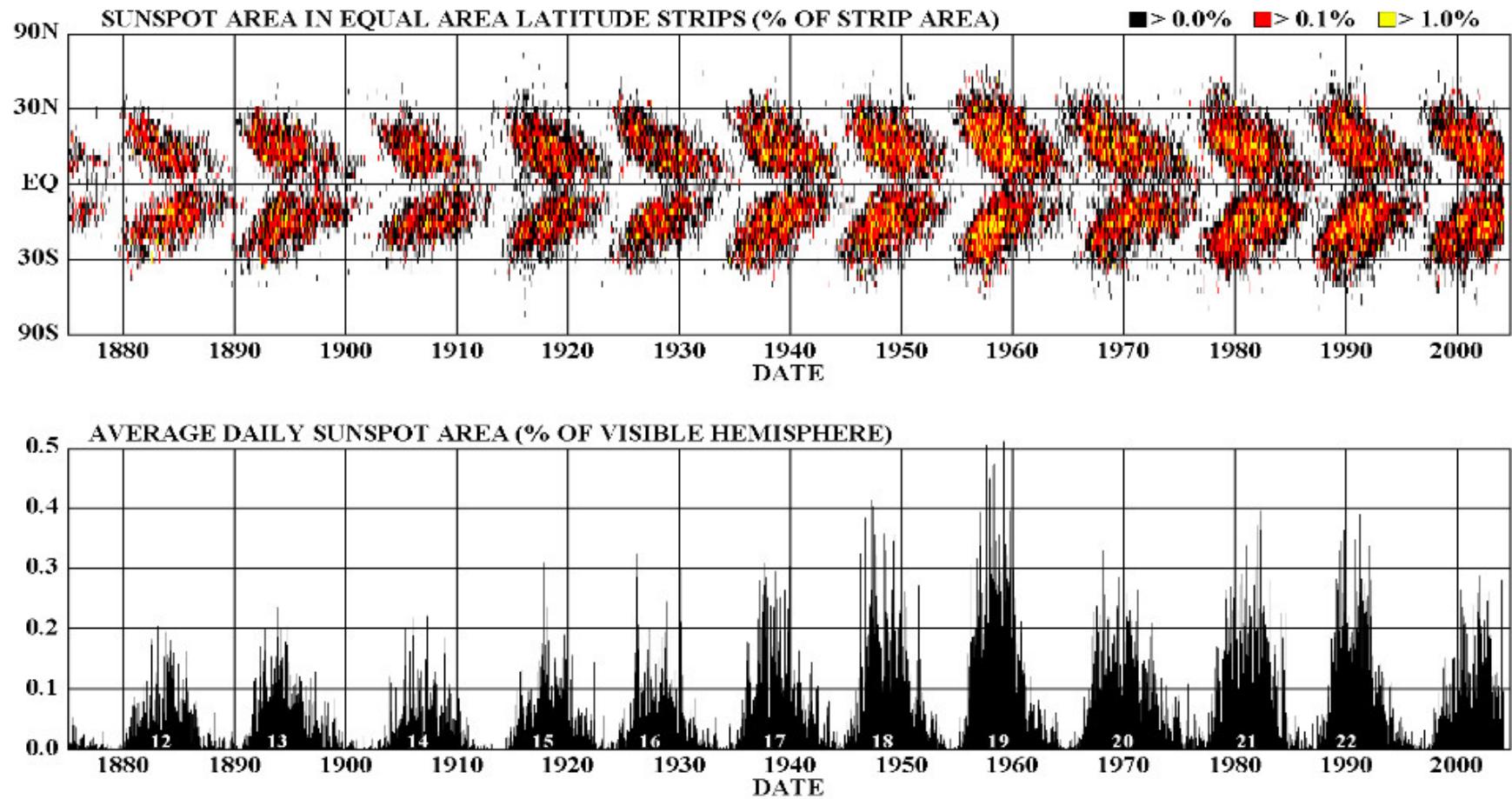
g = # di gruppi di macchie

k fattore correttivo

Il periodo che intercorre tra due massimi (o tra due minimi) è di circa **11 anni**

Il diagramma a farfalla

DAILY SUNSPOT AREA AVERAGED OVER INDIVIDUAL SOLAR ROTATIONS

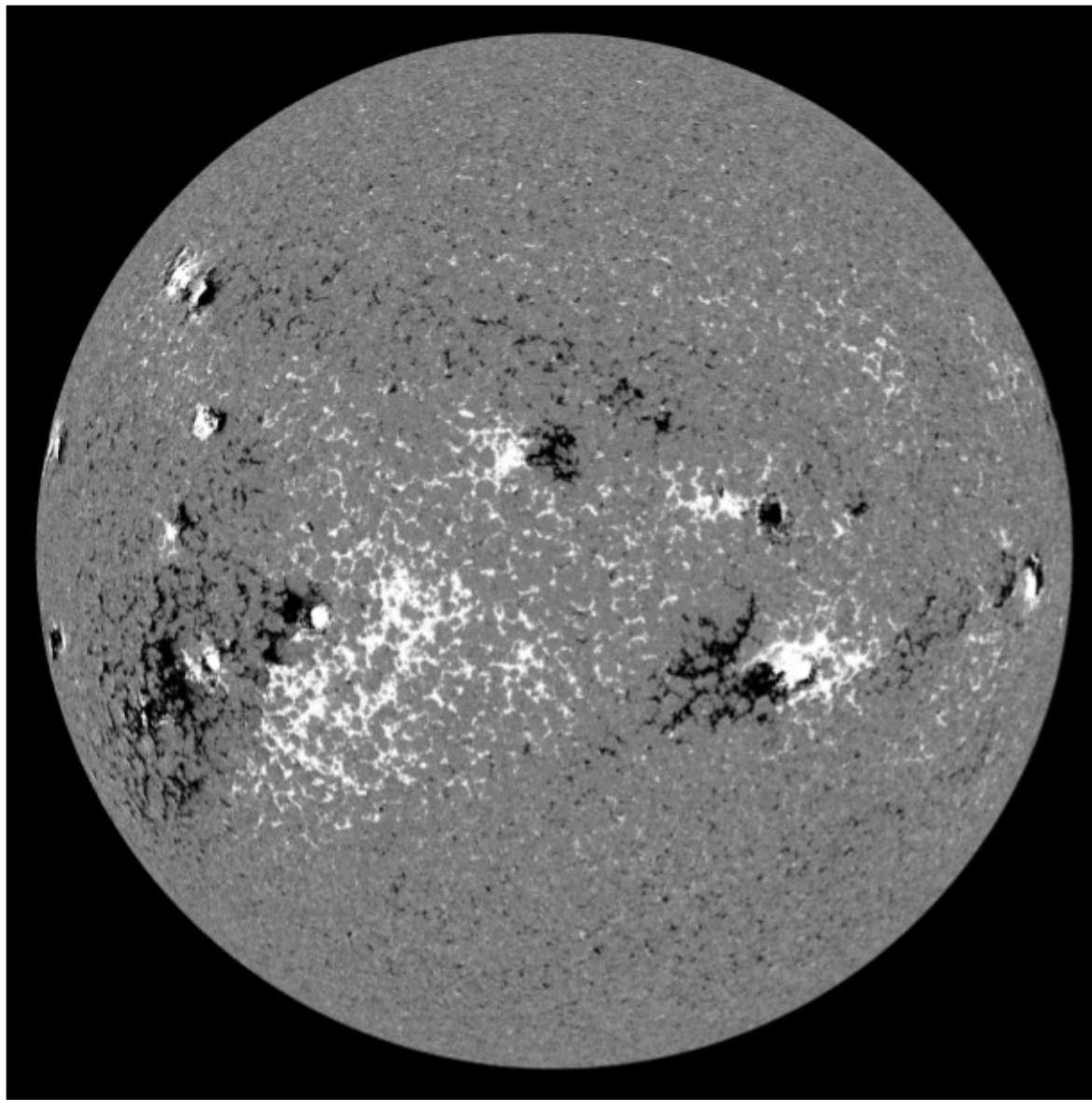


<http://science.msfc.nasa.gov/ssl/pad/solar/images/bfly.gif>

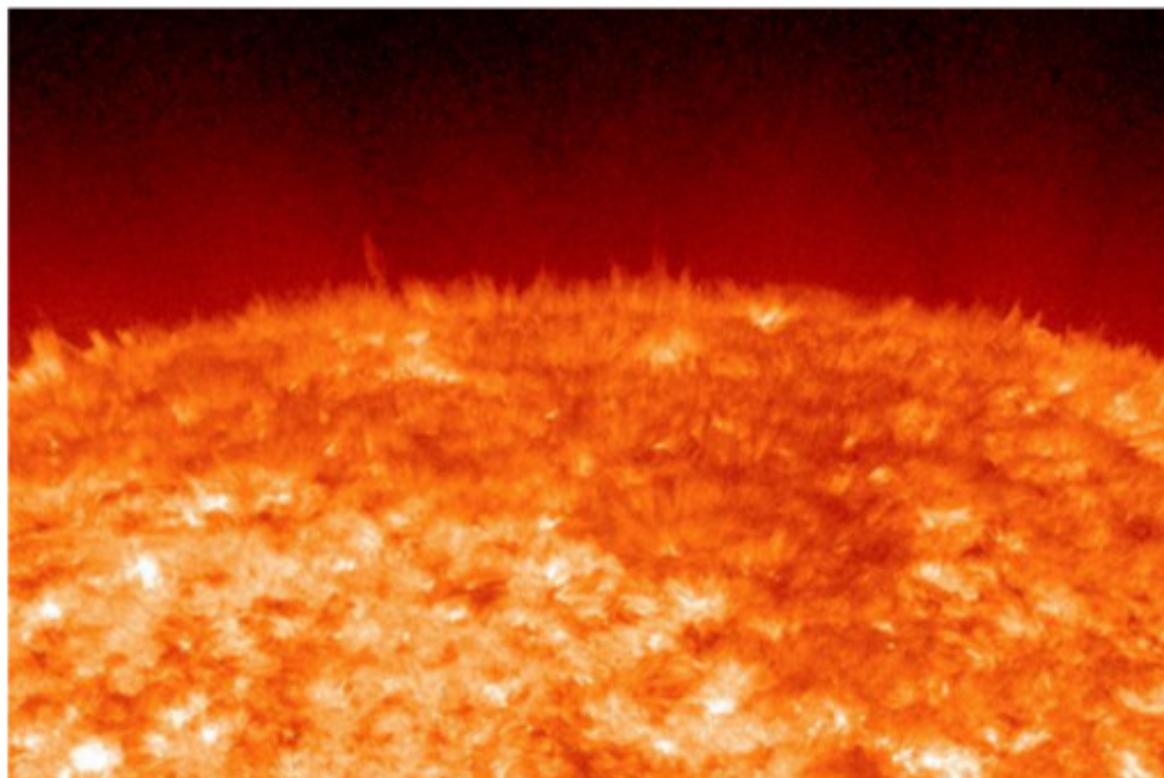
NASA/NSSTC/HATHAWAY 2004/02

Prima della fase di massima attività le macchie sono a latitudini intermedie ($\approx 30^\circ$), quindi 'migrano' verso la zona equatoriale

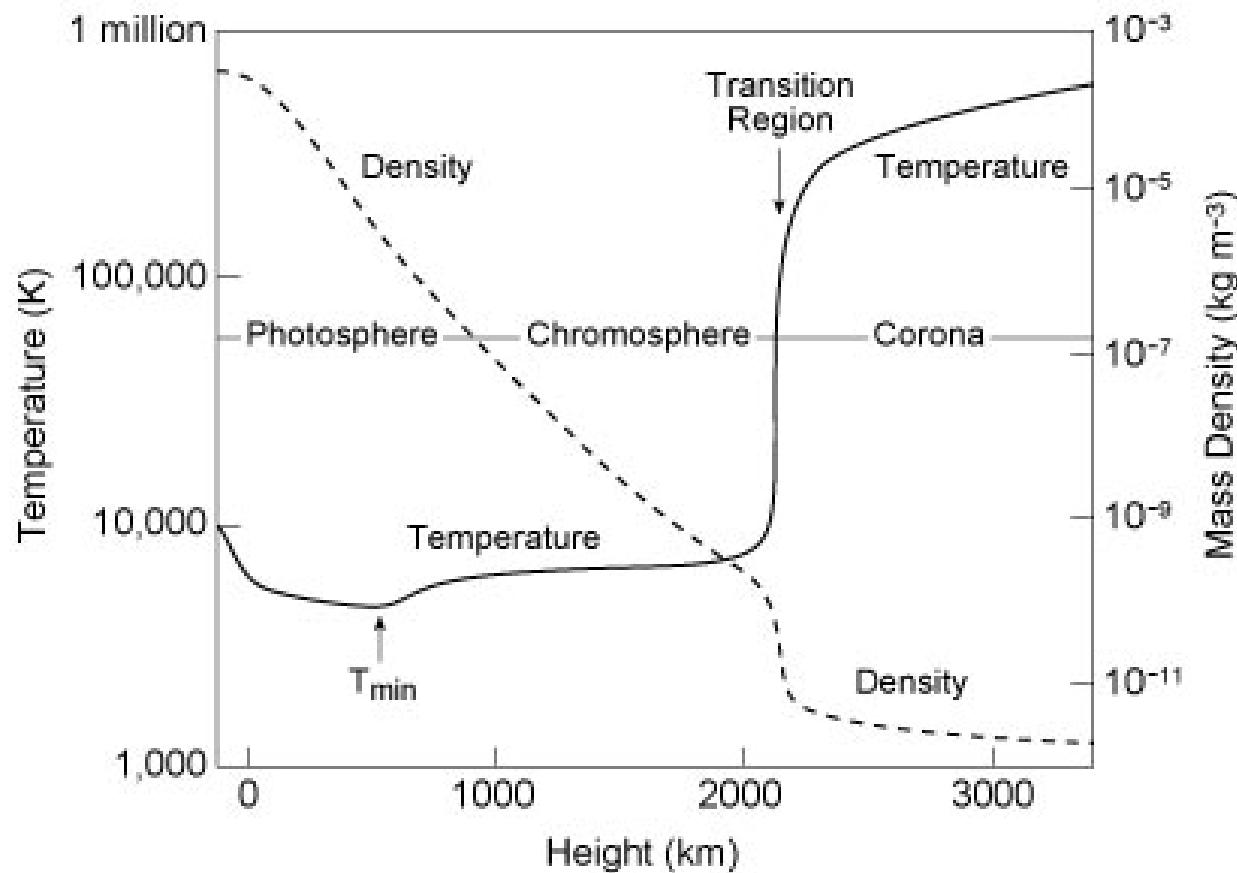
Magnetogramma: effetto Zeeman



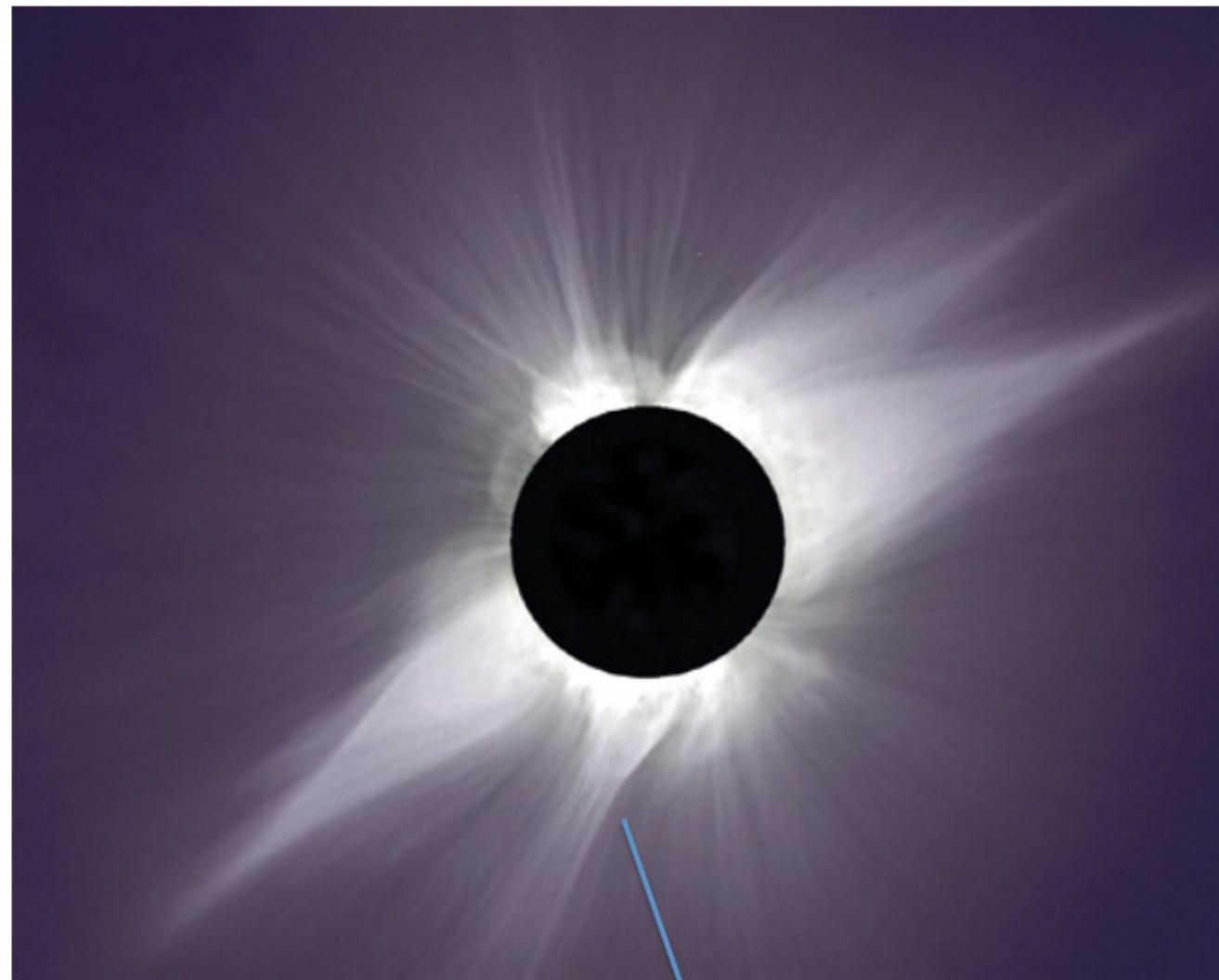
La cromosfera: spicole



La temperatura nell'atmosfera del Sole

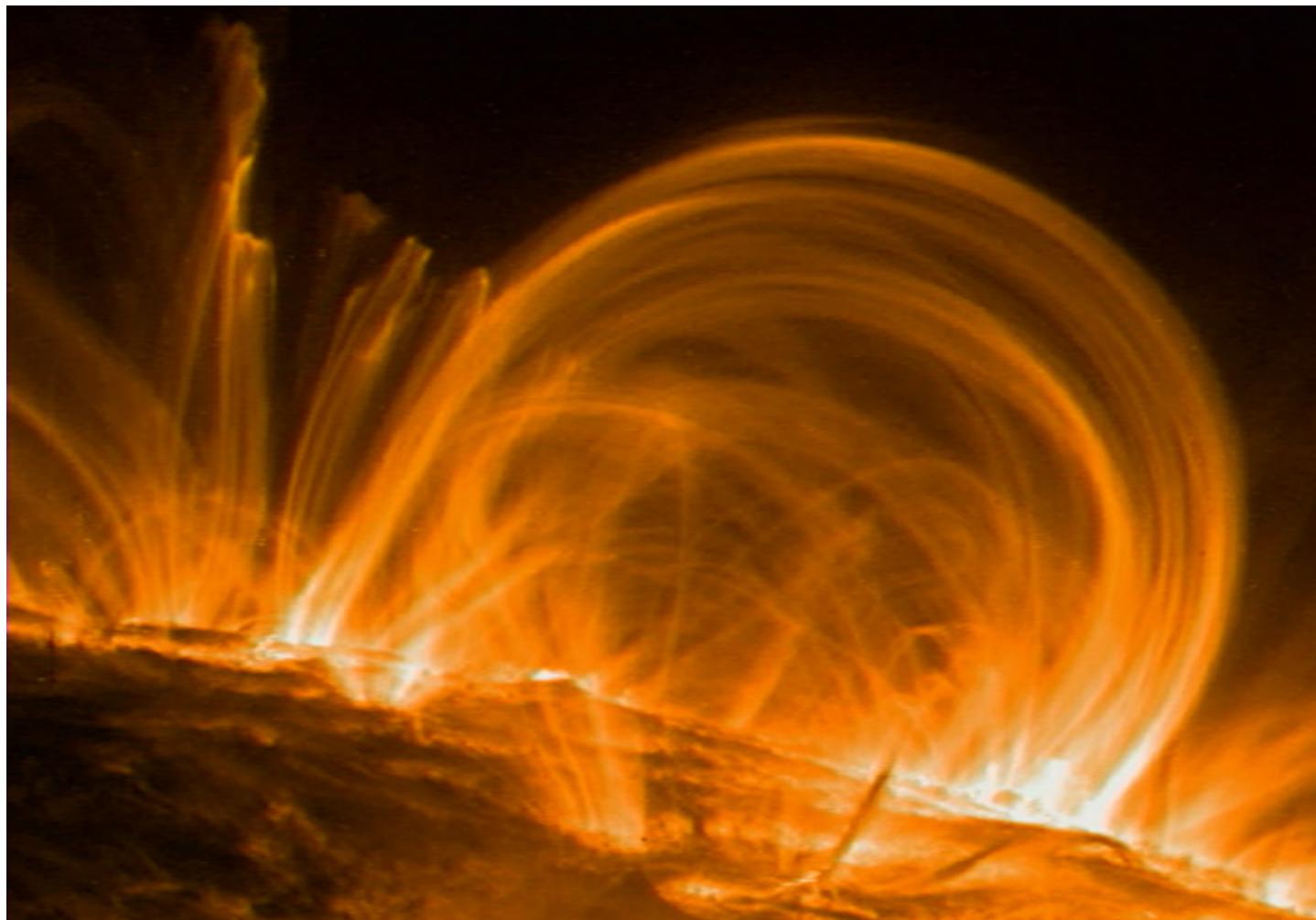


La corona solare

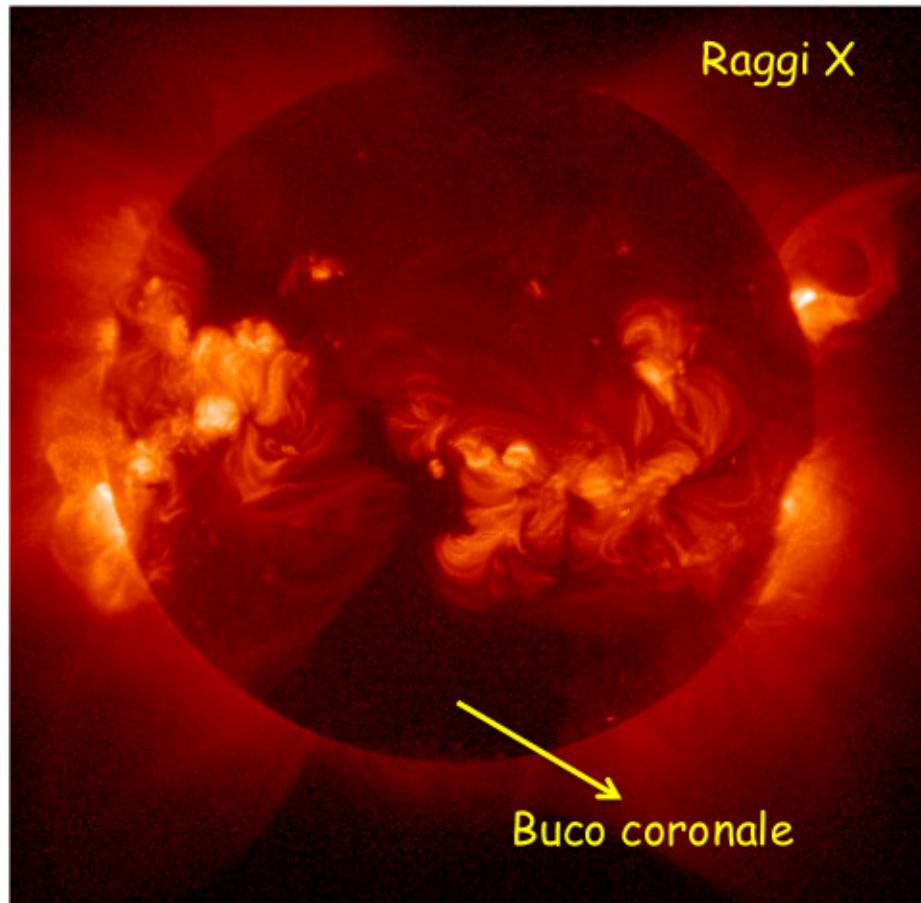


Corona in luce visibile

Archi coronali



La corona nei raggi X



I brillamenti



QUANDO IL SOLE ERUTTA
di Stefano Parisini - Media INAF

INAF TV

Coronal mass ejections

